

PHOTOMETRY OF THE DWARF CEPHEID EH LIBRAE

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Photoelectric B and V observations of the dwarf cepheid EH Lib are given. They enable us to detect small variations in the height of the maxima, whilst the minima appear to be constant. Ten new epochs of maximum are derived. An analysis of the published observations shows that the period is variable. The results are compared to the rate of period change calculated by Dziembowski and Kozłowski.

Key words: dwarf cepheids – AI Velorum stars – photometry

1. INTRODUCTION

EH Lib=BD $-0^{\circ}2911$ is known to be an AI Velorum type variable star since 1950 (Code 1950). On account of the short period EH Lib has been frequently monitored photometrically. The first study, performed photoelectrically by Code, gave some indications that the light curve has fluctuations near the maximum with amplitudes of about 0.05 mag, so Code suspected a possible secondary periodicity effect to be acting in the variable. Since then a notable number of epochs of maximum has been obtained, from photographic or visual measurements, by Ashbrook (1952), Alantsya (1954), Phol (1955), Tsesevich (1956), Batyrev (1952, 1957, 1964), Burnicki and Krygier (1958), Harding and Penston (1966), Braune and Hubscher (1967), Braune *et al.* (1970, 1972), Berdnikov (1972) and Braune and Mundry (1973). Batyrev (1957), taking into account all the previous observations, gave the following elements:

$$\text{Max hel.} = \text{JD } 2433438.6100 + 0.08841320 n.$$

Batyrev (1964) moreover pointed out that the brightness at the maximum was varying and he believed a beat phenomenon was present in the light curves of EH Lib. However the photoelectric measurements obtained by Fitch (1957) and by Fitch *et al.* (1966) give evidence of only small fluctuations of the heights of maxima (less than ± 0.015 mag). Sanwal and Pande (1961) observed EH Lib photoelectrically in two colours along 15 cycles and report that the light curve is quite regular and the period has not changed appreciably since 1950. Oosterhoff and Walraven (1966) obtained five-colour photoelectric measurements, gave an improved period, but did not report about the constancy of the light curve. According to Boardman and Heiser (1972) the period of EH Lib is constant since 1960. Their *uvby* observations however are too few to prove any light curve modulation; a new evaluation of the instant of maximum by least squares cubic fitting is listed in table 1. A photoelectric study performed by Terzan and Rutily (1974) gives new epochs of maximum light, a small correction to the period and a $U-B/B-V$ diagram like that obtained by Oosterhoff and Walraven. The colours however indicate a spectral type at the minimum even later than FO, an unusual value for a 0.09 days period dwarf cepheid. Berdnikov (1975) taking into account 24 instants of maximum, at the most normal epochs, reached the conclusion that the period of EH Lib is varying with the cycle of approximately 1800 days. A radial velocity curve has then been obtained by McNamara and Feltz (1975).

2. LIGHT CURVES AND PERIOD

On account of the contrasting conclusions about the presence of beat phenomena in EH Lib, we thought it useful to obtain a new set of measurements to permit a definitive statement about the stability of the light curves and to analyse all the epochs of maxima listed in the references reported above to check the trend of the period. The observations were obtained with B and V filters on six nights from April through May 1975, using the 102 cm reflector of the Merate Observatory, by means of a conventional

integrating charge photometer. A rapid positioning of the telescope between variable and comparison star, via digital setting controls and short times of integration (typically 20 sec), allowed the reduction of the effects of variable sky transparency on the differential measurements.

EH Lib was compared to BD $-0^{\circ}2909$, used also by Fitch (1957), who gives for it the values: $V=10^m26$, $B-V=+0^m44$; the check star was BD $-0^{\circ}2903$. The Δm between the two comparison stars changed by 0^m03 in the interval GG 42519–549. As it will appear later from the study of the light curves, BD $-0^{\circ}2909$ stayed constant but the check star became brighter.

The light curves of EH Lib are displayed in figure 3 and the individual observations, altogether 900, corrected for light time and differential extinction, are listed in tables 3 and 4. The corrections for the B measurements reach 0^m01 around the maxima, where the colour difference between variable and comparison star is stronger, but elsewhere and for the V measurements they are only a few thousandths of a magnitude or negligible.

No consistent changes are evident in the shapes of the light curves, like those well known in the dwarf cepheids with a Blazhko effect. However to see if EH Lib undergoes small cycle to cycle variations in the brightness at maximum light also during our observing season, as already claimed by Fitch (1957), cubic parabolas were fitted by least squares through the measurements encompassing the maxima and the minima. The epochs and the magnitudes derived are listed in table 1. The mean values are:

$$\begin{aligned} \text{Max } V &= 9^m549 & B-V &= +0^m175 \\ \text{Min } V &= 10.080 & B-V &= +0.312 \\ & & (\text{Max-Min})/\text{Period} &= 0.316 \end{aligned}$$

The uncertainty of a single Δm estimated by the fitting is at the minimum 0^m004 and at the maximum 0^m007 . The latter larger value is likely due to the difficulty of monitoring in two colours the rapid light variations around the maximum light by means of discrete measurements. An inspection of the data in table 1 discloses that the magnitudes at the minimum are constant also if the r.m.e. of a minimum appears to be larger than expected, taking into account the uncertainty quoted above for a single observation. On the contrary the brightness at the maximum shows a systematic trend, as can be seen in figure 1. Representing the magnitudes by means of a linear regression it can be guessed that the maxima decrease 0^m033 each 100 cycles, in both colours.

The course of the period can be examined over an interval of more than one hundred thousand cycles, about 25 years, well covered by the 129 epochs listed in table 2 along with the corresponding sources. The material is quite inhomogeneous: for the visual epochs observed simultaneously by most observers the uncertainty amounts to $\pm 0^d006$; for our ten photoelectric epochs we estimate an internal precision of 0^d0003 (from the fitting); we obtain the same precision for the mean epochs of B and V values. After suitable weights (w) were given, typically $\sqrt{w}=1$ to a visual or photographic individual epoch and $\sqrt{w}=6$ for a photoelectric one, the following ephemerides and their standard errors were calculated:

$$\begin{aligned} \text{Max} = \text{hel. JD } & 2433438.6082 + 0.0884132445 n \\ & \pm \quad \quad \quad 2 \quad \quad \quad 30 \text{ m.e.} \end{aligned}$$

The $O-C$ are listed in table 2. Their distribution does not appear to be gaussian, but a systematic trend results. Moreover we have seen no evidence of periodic oscillation in the residuals, in particular with a period near 1800 days (Berdnikov 1975). Since the dispersion of the visual and photographic residuals, when plotted against the cycle number n , indicates that these epochs have an uncertainty even larger than the value estimated above and more conspicuous than the small variation of the period displayed by the photoelectric observations, it seemed worthwhile studying the period considering the photoelectric epochs only. These observations, plotted in figure 2, appear to be gathered in some groups and give evidence that the period has a complicated variation that we merely represent by a broken straight line. In this way, it gives:

Interval	Period
0– 41000 cycles	0 ^d 088413243 ± 2 m.e.
40000– 45000 cycles	.088413197 31
45000– 78000 cycles	.088413318 2
78000–103000 cycles	.088413142 2

Looking at the m.e. the period seems certainly to have undergone some small long time scale variations.

CONCLUSIONS

The observations we obtained suggest the following conclusions:

a) the brightness of EH Lib seems to change at the maxima, whilst at the minima it remains perceptibly constant. Also if the dispersion of the measurements is rather appreciable in comparison to the effect guessed, the variation of the maxima is real in our opinion and supports the findings of Fitch (1957). Since the light amplitudes of EH Lib measured by the previous observers are about the same as those we observed, it is likely that the maxima change cyclically by a small amount. Our measurements are too few to indicate more about this variation. If the phenomenon should be periodic, a period of about 300 cycles (27 days) can be proposed. On account of the small amount of fluctuation in the maxima, no measurable effect can be expected in the instants of maximum as occurs for the RR stars known to have a Blazhko effect.

b) According to Frolov (Kukarkin 1975) instabilities of the period and of the light curves, as well as the Blazhko effect, are common in the dwarf cepheids. In this respect our results bear out that EH Lib confirms the rule. Dziembowski and Kozłowski (1974) calculate a series of models of low mass stars in the pulsation instability strip and find that the periods should be changed. The rate of change $-d \ln P/dt$ (expressed in year⁻¹) should be comprised between $3 \cdot 10^{-8}$ and $5 \cdot 10^{-7}$. Making a comparison with the values observed for the AI Vel stars they find as possible candidates CY Aqr and EH Lib. In particular the supposed constant period of EH Lib restricts the possible rate of period change which has been guessed. As figure 2 shows, during the last sixty thousand cycles the period of EH Lib becomes shorter and a mean rate of change equal to $3 \cdot 10^{-7}$ can be estimated. However, as the theory supposes a continuous decrease in the period and since it is possible that accidental variations add up to secular changes, only further observations covering a much longer time interval can prove the existence of a secular trend. At present we can only maintain that the period of EH Lib also doesn't appear to be constant, as it occurs for most dwarf cepheids.

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REFERENCES

- Alantsya, I.F.: 1954, *Astron. Circ. USSR* no. 146, 14.
 Ashbrook, J.: 1952, *Astron. J.* **57**, 64.
 Batyrev, A.A.: 1952, *Variable Stars* **8**, 308.
 Batyrev, A.A.: 1957, *Variable Stars* **12**, 137.
 Batyrev, A.A.: 1964, *Variable Stars* **15**, 278.
 Berdnikov, L.N.: 1972, *Variable Stars Suppl.* **1**, 387.
 Berdnikov, L.N.: 1975, *Variable Stars Suppl.* **2**, 199.
 Boardman, W.J. and Heiser, A.M.: 1972, *Publ. Astron. Soc. Pacific* **84**, 680.
 Braune, W. and Hübscher, J.: 1967, *Astron. Nachr.* **290**, 105.
 Braune, W., Hübscher, J. and Mundry, E.: 1970, *Astron. Nachr.* **292**, 185.

- Braune, W., Hübscher, J. and Mundry, E.: 1972, *Astron. Nachr.* **294**, 123.
 Braune, W. and Mundry, E.: 1973, *Astron. Nachr.* **294**, 225.
 Burnicki, A. and Krygier, B.: 1958, *Acta Astron.* **8**, 21.
 Code, A.D.: 1950, *Publ. Astron. Soc. Pacific* **162**, 166.
 Dziembowski, W. and Kozłowski, M.: 1974, *Acta Astron.* **24**, 245.
 Fitch, W.S.: 1957, *Astron. J.* **62**, 108.
 Fitch, W.S., Wisniewski, W.Z. and Johnson, H.L.: 1966, *Arizona Comm.* **71**, 146.
 Harding, G.A. and Penston, M.J.: 1966, *Greenwich Bull.* **115**, 295.
 Kukarkin, B.V.: 1975, *Pulsating Stars*, Wiley and Sons 223.
 McNamara, D.H. and Feltz, K.A.: 1976, *Publ. Astron. Soc. Pacific* **88**, 164.
 Oosterhoff, P.Th. and Walraven, Th.: 1966, *Bull. Astron. Inst. Neth.* **18**, 387.
 Pohl, E.: 1955, *Astron. Nachr.* **282**, 235.
 Sanwal, N.B. and Pande, M.C.: 1961, *Observatory* **81**, 199.
 Terzan, A. and Rutily, B.: 1974, *Astron. Astrophys. Suppl.* **16**, 155.
 Tsesevich, V.P.: 1956, *Astron. Circ. USSR* no. 170, 13.

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Table 1 Light curve parameters of EH Librae

n	Maxima			Minima		
	Hel.J.D.2442	B	V	Hel.J.D.2442	B	V
1	515.5531 ± .0001	9 ^m .707	9 ^m .541	515.5265 ± .0005	10 ^m .385	10 ^m .081
2	515.6410 ± .0005	9.694	9.534	515.6138 ± .0005	10.391	10.069
46	519.5314 ± .0001	9.719	9.539	519.5021 ± .0011	10.387	10.082
47	519.6201 0	9.717	9.559	519.5933 ± .0011	10.401	10.077
69	521.5654 ± .0003	9.727	9.557	521.5386 ± .0007	10.384	10.077
70	—	—	—	521.6229 ± .0010	10.390	10.075
361	547.3819 ± .0002	9.724	9.544	—	—	—
362	—	—	—	547.4437 0	—	10.080
384	549.4154 ± .0002	9.730	9.543	—	—	—
385	549.5032 ± .0005	9.744	9.557	549.4762 ± .0002	10.401	10.091
407	551.4492 ± .0004	9.738	9.559	551.4220 ± .0002	10.393	10.081
408	551.5375 ± .0002	9.737	9.568	551.5101 ± .0008	10.395	10.086

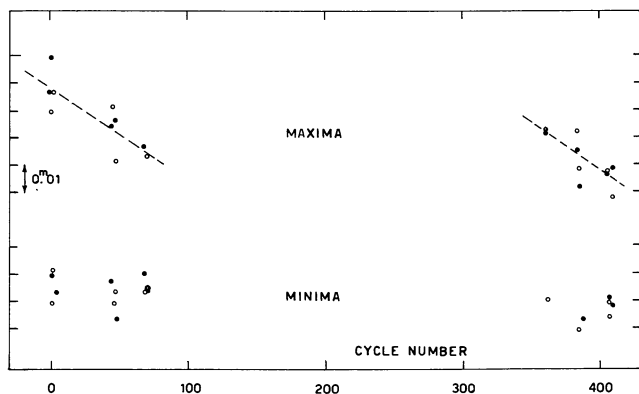


Figure 1 Magnitudes at maximum (top) and minimum light (bottom) of EH Lib. The regression coefficients of the linear regressions give respectively 0.81 and 0.74 for the two groups of maxima. The V measurements (◐) are shifted vertically with respect to the B ones (.) until the two groups merge together.

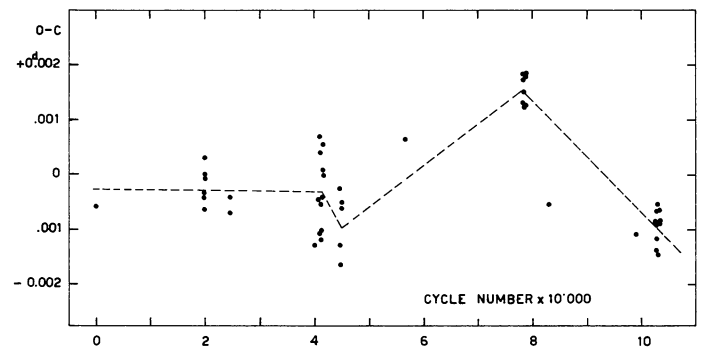


Figure 2 Plot showing variability of period. The residuals are calculated according to the ephemerides: Max = JD 2433438.6082 + 0.0884132445 n.

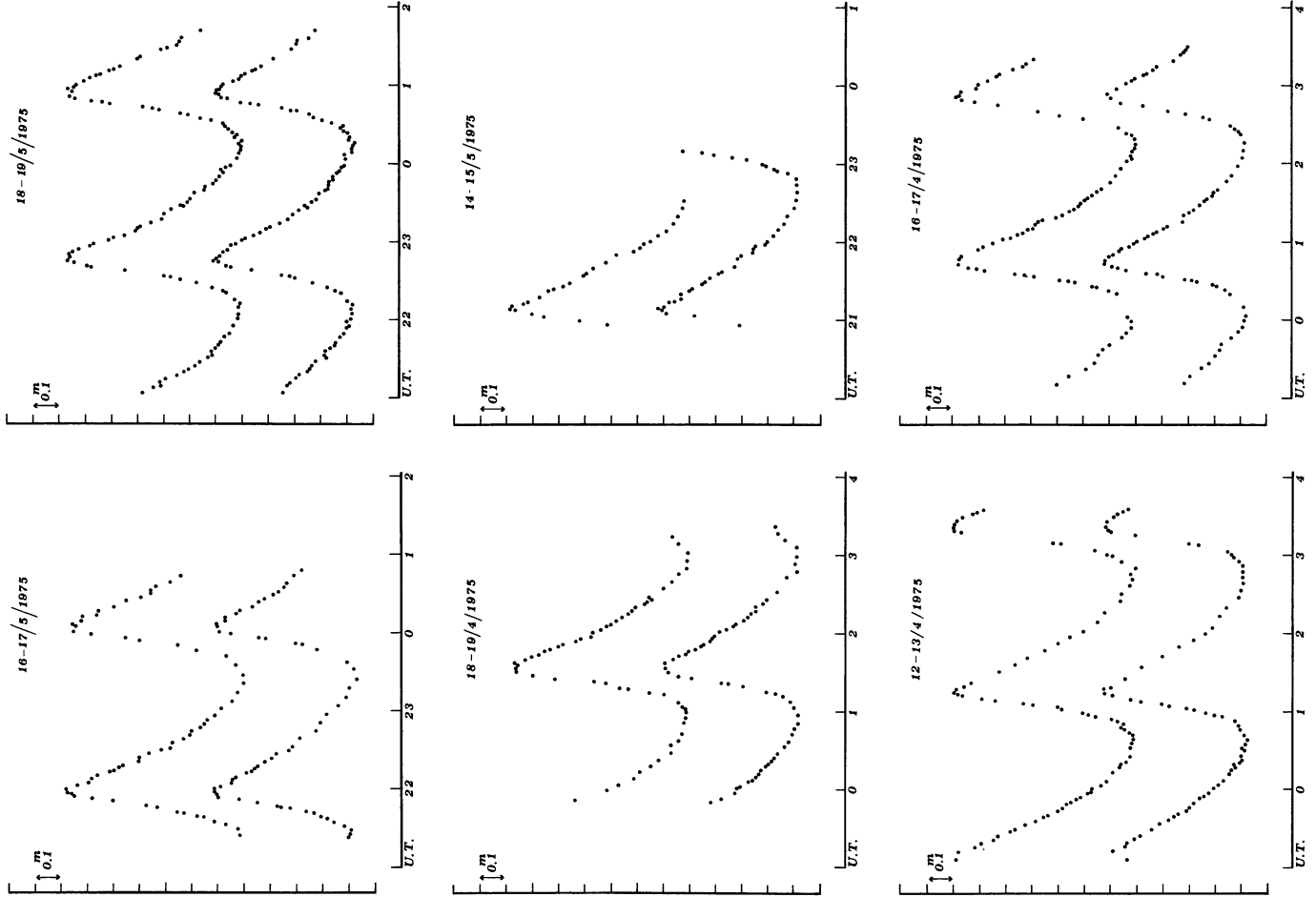


Figure 3 B (upper) and V light curves of EH Librae.

Table 4 V magnitudes of EH Librae

HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V	HEL. J.D. 24.....	V
42515+4629	9.663	42515+5137	9.849	42519+4908	10.025	42519+5742	10.011	42521+5561	10.045
4475	9.566	4511	9.782	4436	10.054	4529	10.053	4526	10.064
4717	9.662	4576	9.675	4500	10.075	4570	10.068	4585	10.076
4739	9.681	4585	9.634	4527	10.342	4598	10.073	4505	10.077
4785	9.718	4590	9.657	4537	10.066	4594	10.060	4570	10.068
4820	9.748	4549	9.539	4542	10.002	4602	10.057	4599	10.019
4850	9.788	4539	9.537	4581	9.967	4619	10.043	4515	9.982
4859	9.803	4595	9.617	4505	9.891	4623	9.946	4543	9.889
4867	9.819	4556	9.681	4513	9.872	4642	9.918	4550	9.865
4881	9.888	4515	9.623	4520	9.749	4638	9.782	4571	9.786
4884	9.888	4548	9.634	4530	9.749	4642	9.782	4571	9.786
4899	9.877	4503	9.878	4554	9.742	4610	9.693	4556	9.669
4936	9.882	4539	9.923	4553	9.700	4618	9.604	4568	9.660
4946	9.902	4529	9.978	4571	9.654	4607	9.555	4564	9.593
4984	9.931	4576	10.014	4585	9.573	4623	9.593	4554	9.567
4990	9.927	4626	10.022	4503	9.543	4653	9.626	4572	9.572
5027	9.970	4606	10.029	4518	9.546	4658	9.661	4568	9.596
5048	9.990	4635	10.063	4534	9.559	4630	9.650	4576	9.619
5056	9.988	4616	10.065	4541	9.562	4634	9.648	4574	9.618
5106	10.011	4621	10.051	4571	9.601	4635	9.708	4573	9.682
5125	10.036	4645	10.068	4579	9.609	4682	9.682	4570	9.705
5132	10.026	4672	10.066	4613	9.656	4624	9.685	4576	9.718
5151	10.040	4606	9.897	4620	9.666	4643	9.685	4583	9.747
5180	10.068	4634	9.861	4538	9.598	4643	9.685	4581	9.727
5209	10.085	4637	9.853	4565	9.728	4652	9.700	4581	9.727
5217	10.073	4677	9.953	4655	9.732	4652	9.733	4585	9.801
5227	10.085	4683	9.955	4672	9.742	4681	9.634	4578	9.630
5270	10.062	4621	9.950	4690	9.745	4616	9.682	4590	9.670
5279	10.075	4655	9.978	4528	9.578	4516	9.562	4590	9.670
5291	10.075	4655	9.978	4528	9.578	4516	9.562	4590	9.670
5300	10.073	4672	9.989	4528	9.578	4516	9.562	4590	9.670
5309	10.075	4672	9.989	4528	9.578	4516	9.562	4590	9.670
5313	10.075	4672	9.989	4528	9.578	4516	9.562	4590	9.670
5331	10.084	4691	9.969	4591	9.885	4589	9.522	4576	9.545
5347	10.047	4651	9.847	4512	9.506	4510	9.534	4594	9.511
5365	10.039	4470	9.507	4520	9.507	4515	9.584	4602	9.586
5373	10.042	4471	9.510	4548	9.534	4515	9.584	4602	9.586
5393	9.983	4412	9.563	4566	9.593	4515	9.584	4602	9.586
5399	9.983	4412	9.563	4566	9.593	4515	9.584	4602	9.586
5409	9.983	4412	9.563	4566	9.593	4515	9.584	4602	9.586
5421	9.976	4421	9.565	4520	9.524	4521	10.026	4605	10.045
42521+4612	9.824	42549+4912	9.840	42549+4912	9.840	42551+4823	10.480	42551+4823	10.480
4634	9.752	4437	9.895	4448	10.480	4448	10.480	4448	10.480
4663	9.630	4471	9.747	4436	10.038	4436	10.038	4436	10.038
42549+3911	10.058	4476	9.719	4436	10.038	4436	10.038	4436	10.038
3351	10.066	4475	9.571	4446	9.588	4446	9.588	4446	9.588
3375	10.044	4439	9.560	4430	9.665	4430	9.665	4430	9.665
3309	9.552	4399	10.011	4399	9.844	4399	9.844	4399	9.844
3323	9.566	4424	9.978	4429	9.904	4429	9.904	4429	9.904
3343	9.582	4438	9.930	4448	9.632	4448	9.620	4448	9.620
3348	9.600	4448	9.898	4453	9.647	4453	9.647	4453	9.647
3387	9.626	4472	9.799	4462	9.718	4462	9.718	4462	9.718
3311	9.664	4477	9.790	4497	9.565	4497	9.565	4497	9.565
3321	9.671	4497	9.710	4507	9.576	4507	9.576	4507	9.576
3325	9.671	4510	9.613	4510	9.613	4510	9.613	4510	9.613
3395	9.721	4431	9.562	4454	9.617	4454	9.617	4454	9.617
3379	9.742	4436	9.561	4569	9.827	4565	9.821	4583	9.847
3397	9.760	4465	9.549	4465	9.549	4465	9.549	4465	9.549
3397	9.760	4465	9.549	4465	9.549	4465	9.549	4465	9.549
4116	9.905	4480	9.572	42551+3779	9.814	4480	9.572	4480	9.572
4136	9.910	4499	9.613	3818	9.833	4419	9.077	4537	9.457
4136	9.910	4499	9.613	3818	9.833	4419	9.077	4537	9.457
4204	9.782	4434	9.628	3862	9.869	4463	9.754	4552	9.581
4133	9.832	4453	9.670	3877	9.878	4463	9.754	4552	9.581
4482	9.847	4482	9.847	3866	9.912	4483	9.806	4536	9.597
4465	9.955	4487	9.713	3931	9.928	4471	9.857	4545	9.579
4416	9.985	4431	9.788	3931	9.928	4471	9.857	4545	9.579
4199	9.985	4431	9.788	3931	9.928	4471	9.857	4545	9.579
4307	10.026	4455	9.788	3984	9.975	4475	9.901	4549	9.627
4307	10.026	4455	9.788	3984	9.975	4475	9.901	4549	9.627
4334	10.075	4482	9.936	4482	9.936	4482	9.936	4482	9.936
4331	10.065	4499	9.847	4419	9.597	4410	9.445	4549	9.627
4478	10.072	4478	9.955	4419	9.597	4410	9.445	4549	9.627
4478	10.072	4478	9.955	4419	9.597	4410	9.445	4549	9.627
4507	10.012	4470	9.972	4488	9.977	4488	9.977	4488	9.977
4511	9.992	4451	9.992	4436	10.066	4432	10.011	4532	10.011
4511	9.992	4451	9.992	4436	10.066	4432	10.011	4532	10.011
4585	9.925	4410	10.083	4460	10.077	4461	10.035	4570	9.875
4585	9.925	4410	10.083	4460	10.077	4461	10.035	4570	9.875
4600	9.877	4449	10.055	4499	10.079	4461	10.035	4570	9.875